

Dynamics between order and chaos in conceptual models of glacial cycles

Takahito Mitsui · Kazuyuki Aihara

Abstract The dynamics of glacial cycles is studied in terms of the dynamical systems theory. We explore the dependence of the climate state on the phase of astronomical forcing by examining five conceptual models of glacial cycles proposed in the literature. The models can be expressed as quasiperiodically forced dynamical systems. It is shown that four of them exhibit a strange nonchaotic attractor (SNA), which is an intermediate regime between quasiperiodicity and chaos. Then, the dependence of the climate state on the phase of astronomical forcing is not given by smooth relations, but constitutes a geometrically strange set. Our result suggests that the dynamics of SNA is a candidate for the dynamics of glacial cycles, in addition to the well-known dynamics of quasiperiodicity and chaos.

Keywords Glacial cycles · Astronomical theory of climate change · Conceptual models · Quasiperiodically forced dynamical systems · Strange nonchaotic attractors

1 Introduction

The climate change in the Quaternary is characterized by alternations between cold (glacial) and warm (interglacial) periods, which are called glacial-interglacial

T. Mitsui
FIRST, Aihara Innovative Mathematical Modelling Project,
JST, 4-6-1 Komaba, Meguro-ku, Tokyo 153-8505, Japan
Institute of Industrial Science, University of Tokyo, 4-6-1
Komaba, Meguro-ku, Tokyo 153-8505, Japan
Tel.: +81-3-5452-6370
Fax: +81-3-5452-6371
E-mail: mitsui@sat.t.u-tokyo.ac.jp

K. Aihara
Institute of Industrial Science, University of Tokyo, 4-6-1
Komaba, Meguro-ku, Tokyo 153-8505, Japan

cycles or alternatively glacial cycles. The aim of this article is to explore the dynamics of glacial cycles by examining conceptual models proposed in the literature.

Astronomical theories of ice ages, as typified by the Milankovitch theory (1941), attempt to explain glacial cycles based on the change in the incoming solar radiation (insolation) due to the long-term variations of the Earth's orbital parameters. Assuming a constant solar output and a perfectly transparent atmosphere, the insolation F at a given latitude and season is a function of the *astro-insolation parameters*: the eccentricity e of the Earth's orbit, the obliquity ε (the inclination of the equator on the ecliptic), and the climatic precession $e \sin \tilde{\omega}$, where $\tilde{\omega}$ is the longitude of the perihelion measured from the moving vernal equinox (Berger 1978). That is, $F = F(e, \varepsilon, e \sin \tilde{\omega})$.

A motion expressed by a superposition of periodic motions of two or more incommensurate frequencies is called a *quasiperiodic motion*. In the studies of the Quaternary climate changes, the astro-insolation parameters are often assumed to be quasiperiodic (Berger 1978):

$$\begin{aligned} e &= e_0 + \sum E_i \cos(\lambda_i t + \phi_i), \\ \varepsilon &= \varepsilon^* + \sum A_i \cos(f_i t + \delta_i), \\ e \sin \tilde{\omega} &= \sum P_i \sin(\alpha_i t + \zeta_i). \end{aligned}$$

Then, the insolation F is also quasiperiodic in time (cf. Eq. (5)). The insolation variation $F(t)$ is known as *astronomical forcing*. The primary frequencies of astronomical forcing are approximately $1/19 \text{ kyr}^{-1}$ and $1/23 \text{ kyr}^{-1}$ for the climatic precession, and approximately $1/41 \text{ kyr}^{-1}$ for the obliquity. The assumption of quasiperiodicity of astronomical forcing is considered to be approximately valid over the last several million years (Berger and Loutre 1991) but may not be valid beyond this period because of the intrinsic

chaoticity of planetary motion (Lasker 1989). However, since our study focuses on the glacial cycles during the past 0.7 Myr after the mid-Pleistocene transition (Clark et al. 2006), we construct our theory on the assumption of quasiperiodic astronomical forcing.

Hays et al. (1976) presented strong evidence for astronomical theories of ice ages. They found the primary frequencies of astronomical forcing in the geological spectra of marine sediment cores. However, the dominant frequency in geological spectra is approximately $1/100 \text{ kyr}^{-1}$, although this frequency component is negligible in the astronomical forcing. This is referred to as the “100 kyr problem.”

To understand the climate responses to astronomical forcing, several concepts related to the dynamical systems theory have been proposed in the literature. The *linear response* is a simple framework to explain the 41, 23, and 19 kyr periodicities in the geological spectra. However, the linear response cannot appropriately account for the 100 kyr periodicity (Hays et al. 1976). Ghil (1994) explained the appearance of the 100 kyr periodicity as a *nonlinear resonance* to the combination tone $1/109 \text{ kyr}^{-1}$ between precessional frequencies $1/19 \text{ kyr}^{-1}$ and $1/23 \text{ kyr}^{-1}$. Contrary to the linear resonance, the nonlinear resonance can occur even if the forcing frequencies are far from the internal frequency of the response system. Benzi et al. (1982) proposed *stochastic resonance* as a mechanism of the 100 kyr periodicity, where the response to a small external forcing is amplified by the effect of noise. See Ganopolski and Rahmstorf (2002) and Braun et al. (2007) for the recent progress related to stochastic resonance. Tziperman et al. (2006) proposed that the timing of deglaciations is set by the astronomical forcing via the *phase-locking* mechanism. If a phase locking occurs, the condition $qf = p_1f_1 + p_2f_2 + \dots + p_Nf_N$ is satisfied, where f is the response frequency, f_1, f_2, \dots, f_N are forcing frequencies, and p_1, p_2, \dots, p_N , and $q(\neq 0)$ are integers (Romeiras et al. 1987). De Saedeleer et al. (2013) suggested *generalized synchronization* (GS) to describe the relation between the glacial cycles and the astronomical forcing. GS means that there is a functional relation between the climate state and the state of astronomical forcing. They also showed that the functional relation may not be unique for a certain model. However, the nature of the relation remains to be elucidated.

In this research, we analyze five conceptual models of glacial cycles proposed in the literature and show that the dependence of the climate state on the phase of astronomical forcing may not be given by smooth relations, but constitutes some geometrically strange set known as a *strange nonchaotic attractor* (SNA).

The remainder of this article is organized as follows. In Section 2, we review the notions of quasiperiodically forced systems and attractors. In Section 3, five models of glacial cycles are explained. In Section 4, we analyze the models numerically and show possible relations between the climate state and the phase of astronomical forcing. In Section 5, we mention a dynamical property of climate models with SNAs. Section 6 concludes this article.

2 Quasiperiodically forced systems and attractors

2.1 Quasiperiodically forced systems

We consider the dynamics of glacial cycles in the framework of *quasiperiodically forced dynamical systems*. Let $\mathbb{T}^N = \mathbb{R}^N / (2\pi\mathbb{Z})^N$ be an N -dimensional torus. In general, quasiperiodically forced dynamical systems can be expressed as

$$\begin{aligned}\dot{\theta} &= \omega, & \theta \in \mathbb{T}^N, \\ \dot{\mathbf{x}} &= \mathbf{g}(\mathbf{x}, \theta), & \mathbf{x} \in \mathbb{R}^M,\end{aligned}\tag{1}$$

where $\theta = (\theta_1, \theta_2, \dots, \theta_N)^T$ is the phase of the drive subsystem, $\mathbf{x} = (x_1, x_2, \dots, x_M)^T$ is the state of the response subsystem, $\mathbf{g}(\mathbf{x}, \theta)$ a periodic function in each phase θ_i , and $\omega = (\omega_1, \omega_2, \dots, \omega_N)^T$ is a vector of incommensurate frequencies such that $k_1\omega_1 + k_2\omega_2 + \dots + k_N\omega_N = 0$ does not hold for any set of integers, k_1, k_2, \dots, k_N , except for the trivial solution $k_1 = k_2 = \dots = k_N = 0$. In the astronomical theory of climate change, θ corresponds to the phase of astronomical forcing $F(t)$, and \mathbf{x} corresponds to the climate state, as described in Section 3.

2.2 Attractors

Like most conceptual models of glacial cycles, we assume the existence of *attractors* in systems (1). An attractor is a compact set with a neighborhood such that, for almost every initial condition in this neighborhood, the limit set of the orbit as time tends to infinity is the attractor (Grebogi et al. 1984). The closure of set of initial conditions which approach an attractor as time tends to infinity is called the *basin of attraction* of the attractor. Note that the attractor of system (1) is a subset of $\mathbb{T}^N \times \mathbb{R}^M$. Thus, the dependence of the climate state on the phase of astronomical forcing is represented by this attractor, after a certain transient time. In quasiperiodically forced systems (1), the following types of attractors typically appear: *N-frequency*

quasiperiodic attractors, L-frequency quasiperiodic attractors with $L > N$, Strange nonchaotic attractors (SNAs) and chaotic attractors.

Lyapunov exponents give the mean exponential rates of divergence of nearby orbits. Although the Lyapunov exponents depend on initial conditions, they have the same set of values for almost every initial condition, under general circumstances. Such initial conditions or orbits which give the same set of values of the Lyapunov exponents are called *typical*. A chaotic attractor is an attractor for which typical orbits on the attractor have a positive Lyapunov exponent (Grebogi et al. 1984). That is, almost all orbits on the attractor have sensitive dependence on initial conditions. In quasiperiodically forced systems (1), an attractor is chaotic if the maximal conditional Lyapunov exponent (CLE) λ is positive for typical orbits on the attractor, and non-chaotic otherwise. The maximal CLE λ is defined as

$$\lambda = \lim_{t \rightarrow \infty} \frac{1}{t} \ln \frac{|\delta \mathbf{x}(t)|}{|\delta \mathbf{x}(0)|},$$

where $\delta \mathbf{x}(t)$ is an infinitesimal displacement of the orbit from $\mathbf{x}(t)$. The evolution of the infinitesimal displacement is given by the variational equation $\dot{\delta \mathbf{x}} = D_{\mathbf{x}} G(\mathbf{x}, \theta) \delta \mathbf{x}$.

An N -frequency quasiperiodic attractor is an attracting N -dimensional torus on which the motion is quasiperiodic. N -frequency quasiperiodic attractors are characterized by a negative maximal CLE, $\lambda < 0$.

An L -frequency quasiperiodic attractor with $L > N$ is an attracting L -dimensional torus on which the motion is quasiperiodic. L -frequency quasiperiodic attractors with $L > N$ are characterized by a maximal CLE of zero, $\lambda = 0$.

A strange nonchaotic attractor (SNA) is a strange attractor for which typical orbits have nonpositive Lyapunov exponents (Grebogi et al. 1984). The strange attractor is defined as an attractor which is not a finite set of points and is neither piecewise differentiable curve or surface, nor a volume bounded by a piecewise differentiable closed surface. SNAs are characterized by a negative maximal CLE, $\lambda < 0$. It is proven that some SNAs are represented by a closure of graphs of almost everywhere discontinuous functions (Keller 1996; Alseda and Costa 2009).

SNAs lie in an intermediate regime between quasiperiodicity and chaos. The power spectrum of an SNA is intermediate one between quasiperiodicity and chaos, namely, a singular continuous (fractal) spectrum (Feudel et al. 1996). SNAs are multifractal objects in the sense that the capacity dimension is larger than the information dimension (Ding et al. 1989; Hunt and Ott 2001). For comprehensive reviews of SNAs, refer to Prasad et al. (2001) as well as Feudel et al. (2006).

In the remainder of this article, we will not encounter examples of L -frequency quasiperiodic attractors with $L > N$, and sometimes N -frequency quasiperiodic attractors are simply referred to as quasiperiodic attractors.

2.3 Examples of N -frequency quasiperiodic attractors and an SNA

A quasiperiodically forced map is regarded as a stroboscopic map obtained by sampling the state of the continuous-time system (1) at time intervals corresponding to one of the forcing frequencies. Let us consider a quasiperiodically forced map studied by Glendinning (2002):

$$\begin{aligned} \theta(n+1) &= \theta(n) + \omega \pmod{1}, \\ x(n+1) &= 2\sigma(\alpha + \cos 2\pi\theta(n)) \tanh x(n), \end{aligned} \quad (2)$$

where $(\theta, x) \in [0, 1] \times \mathbb{R}$, $n \in \mathbb{Z}$, $\omega \in \mathbb{R} \setminus \mathbb{Q}$, and σ and α are non-negative parameters.

Figure 1(a) shows the quasiperiodic attractors for $\sigma = 1.2$, $\alpha = 1.001$, and $\omega = (\sqrt{5} - 1)/2$ (solid lines). Each attractor is represented by a graph of a smooth function. If we denote the attractor in the region $x > 0$ by $x = \phi(\theta)$, the attractor in the other region is $x = -\phi(\theta)$ because of the symmetry $x \rightarrow -x$. The line $x = 0$ is an unstable invariant set, namely, a *repellor*. In this case, the repellor constitutes the basin boundary of two attractors.

Figure 1(b) shows the SNA for $\sigma = 1.2$, $\alpha = 0.8$, and $\omega = (\sqrt{5} - 1)/2$. The SNA is represented by the closure of two graphs $x = \varphi(\theta)$ and $x = -\varphi(\theta)$ having following properties (Glendinning 2002; Keller 1996):
(i) $\varphi(\theta) \geq 0$ for all $\theta \in [0, 1]$;
(ii) the union of two graphs $x = \varphi(\theta)$ and $x = -\varphi(\theta)$ is invariant under the map;
(iii) typical conditional Lyapunov exponents on the graphs are negative;
(iv) $\varphi(\theta)$ is discontinuous at almost all $\theta \in [0, 1]$;
(v) $\varphi(\theta)$ is upper semi-continuous;
(vi) $\varphi(\theta) = 0$ on a dense, but measure zero, set of values of $\theta \in [0, 1]$.

Unlike quasiperiodic attractors, SNAs contain some repellors in themselves (Stark 1997). See the repellor $x = 0$ embedded in the SNA in Fig. 1(b). Due to such repellors, the orbits on SNAs become temporarily unstable and sensitive to perturbations, as shown in Section 5.

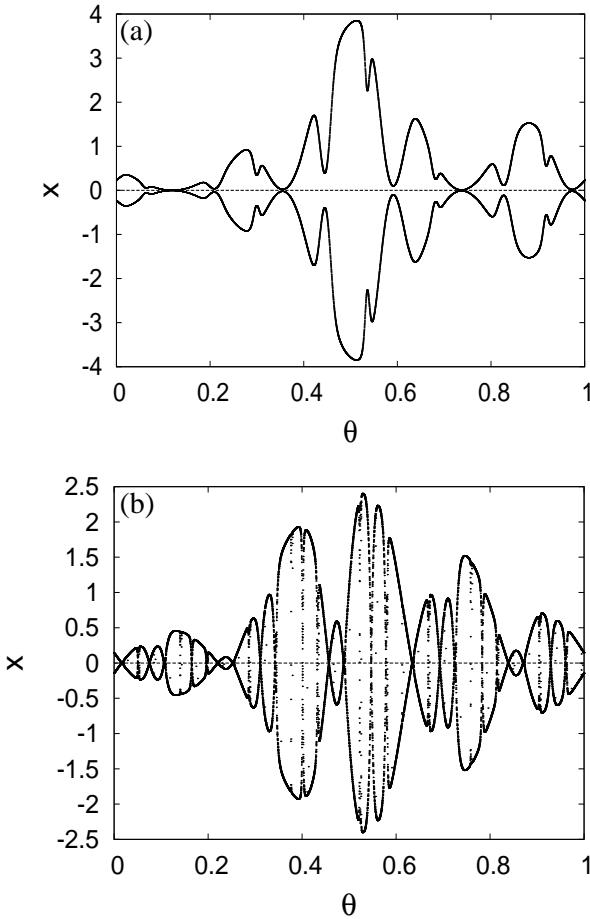


Fig. 1 Attractors in map (2): **a** Quasiperiodic attractors for $\sigma = 1.2$, $\alpha = 1.001$, and $\omega = (\sqrt{5}-1)/2$ (upper and lower solid lines). The maximal CLE is $\lambda = -0.3677$ for each attractor. **b** SNA for $\sigma = 1.2$, $\alpha = 0.8$, and $\omega = (\sqrt{5}-1)/2$ (dots). The maximal CLE is $\lambda = -0.2124$. In each panel, the dashed line $x = 0$ is the unstable invariant set

2.4 Numerical characterization of strangeness of attractors

SNAs and N -frequency quasiperiodic attractors can be numerically distinguished by the phase sensitivity exponent (Pikovsky and Feudel 1995) or by the parameter sensitivity exponent (Nishikawa and Kaneko 1996). First, let us consider the case of the continuous-time system (1). The phase sensitivity function $\Gamma(t)$ of a variable x_j with respect to a variable θ_i is defined as

$$\Gamma(t) = \min_{\{\mathbf{x}(0), \theta(0)\}} \left\{ \max_{0 \leq \tau \leq t} \left| \frac{\partial x_j(\tau)}{\partial \theta_i(0)} \right| \right\}.$$

The derivative $\partial x_j(\tau)/\partial \theta_i(0)$ characterizes the sensitivity of x_j with respect to a change of θ_i . This derivative is obtained by integrating

$$\frac{d}{dt} \frac{\partial x_j}{\partial \theta_i(0)} = \sum_{k=1}^M \frac{\partial g_j(\mathbf{x}, \theta)}{\partial x_k} \frac{\partial x_k}{\partial \theta_i(0)} + \frac{\partial g_j(\mathbf{x}, \theta)}{\partial \theta_i(0)} \quad (3)$$

along the solution $(\theta(t), \mathbf{x}(t))$. For SNAs, the phase sensitivity function grows as $\Gamma(t) \simeq t^\mu$ with $\mu > 0$, where μ is called the *phase sensitivity exponent*. On the other hand, for N -frequency quasiperiodic attractors, $\Gamma(t)$ is bounded and $\mu = 0$.

Assume that function $\mathbf{g}(\mathbf{x}, \theta)$ contains some parameter, say α . The parameter sensitivity function $\Psi(t)$ of variable x_i with respect to parameter α is defined as

$$\Psi(t) = \min_{\{\mathbf{x}(0), \theta(0)\}} \left\{ \max_{0 \leq \tau \leq t} \left| \frac{\partial x_j(\tau)}{\partial \alpha} \right| \right\}.$$

The derivative $\partial x_j(\tau)/\partial \alpha$ characterizes the sensitivity of x_j with respect to a change of α . This derivative is obtained by integrating

$$\frac{d}{dt} \frac{\partial x_j}{\partial \alpha} = \sum_{k=1}^M \frac{\partial g_j(\mathbf{x}, \theta)}{\partial x_k} \frac{\partial x_k}{\partial \alpha} + \frac{\partial g_j(\mathbf{x}, \theta)}{\partial \alpha} \quad (4)$$

along the solution $(\theta(t), \mathbf{x}(t))$. For SNAs, the phase sensitivity function grows as $\Psi(t) \simeq t^\nu$ with $\nu > 0$, where ν is called the *parameter sensitivity exponent*. On the other hand, for N -frequency quasiperiodic attractors, $\Psi(t)$ is bounded and $\nu = 0$.

For discrete-time systems, such as Eq. (2), the methods for characterizing the strangeness of attractors are almost the same except that time evolutions corresponding to Eqs. (3) and (4) become time-discrete (Pikovsky and Feudel 1995; Nishikawa and Kaneko 1996).

Figure 2 shows the behaviors of the phase sensitivity function $\Gamma(n)$ and the parameter sensitivity function $\Psi(n)$ with respect to α for the attractors in Fig. 1. The functions $\Gamma(n)$ and $\Psi(n)$ saturate for large n for the quasiperiodic attractors, and on the other hand, show power-law divergences with exponents $\mu = 1.03$ and $\nu = 0.99$, respectively, for the SNA.

3 Models

3.1 Astronomical insolation forcing

Milankovitch (1941) considered the high-latitude northern hemisphere summer insolation as a decisive factor for glaciations. We use the insolation $F(t)$ at 65°N on the day of the summer solstice in the simulations of all the models. For simplicity in computation, the insolation $F(t)$ is calculated using the formula given in De Saedeleer et al. (2013)¹:

$$F(t) = \frac{1}{a} \sum_{i=1}^N [s_i \sin(\omega_i t) + c_i \cos(\omega_i t)],$$

¹ According to De Saedeleer et al. (2013), the formula is valid over the past 1 Myr, and its mean error compared to Lasker et al. (2004) is 6.7 W/m^2 with peaks at 27.5 W/m^2 .

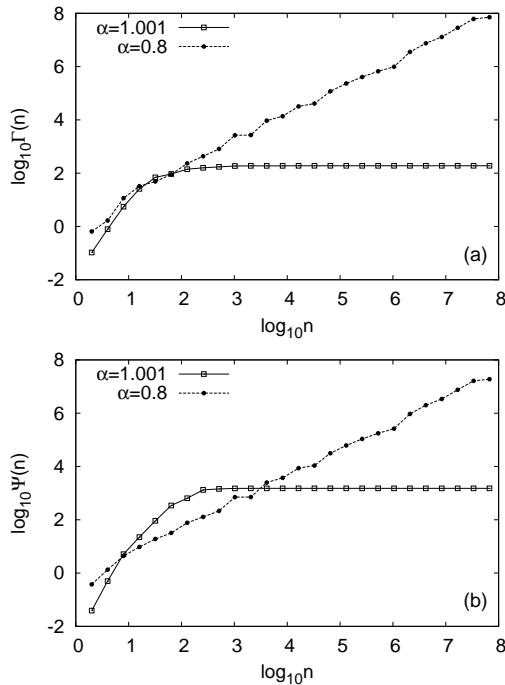


Fig. 2 **a** Phase sensitivity function $\Gamma(n)$ for the quasiperiodic attractors of map (2) with $\sigma = 1.2$, $\alpha = 1.001$, and $\omega = (\sqrt{5}-1)/2$ (open square) and for the SNA with $\sigma = 1.2$, $\alpha = 0.8$, and $\omega = (\sqrt{5}-1)/2$ (filled circle). **b** Phase sensitivity function $\Psi(n)$ for the same attractors

where $N = 35$ is the number of different frequencies ω_i included in the sum. The values of frequencies ω_i are those in Berger (1978). The coefficients s_i and c_i are extracted from the solution of Berger (1978) by performing a linear regression (De Saedeleer et al. 2013). To make the paper self-contained, the values of ω_i , s_i , and c_i are listed in Table 3 in Appendix 1. The parameter a is a scaling constant, the value of which is different among models. The top panel of Fig. 4 shows the insolation $F(t)$ normalized to have a zero-mean and unit variance (with $a = 23.58$).

Introducing variables $\theta_i(t) = \omega_i t$ ($i = 1, 2, \dots, N$), the above equation for $F(t)$ can be written as

$$F(t) = \frac{1}{a} \sum_{i=1}^N [s_i \sin \theta_i(t) + c_i \cos \theta_i(t)], \quad (5)$$

$$\dot{\theta}_i = \omega_i, \quad \theta_i(0) = 0 \quad (i = 1, 2, \dots, N).$$

Note that θ_i are the phase variables of astronomical forcing in climate models.

Table 1 Sets of parameter values used for the SM90 and MS90 models

Set	p	q	r	s	u	v	w
A	1.0	2.5	0.9	1.0	0.6	0.2	0.5
B	0.82	4.24	0.95	0.53	0.32	0.02	0.66
C	1.0	1.2	0.8	0.8	0.7	0.0	0.0

3.2 Conceptual models of glacial cycles

3.2.1 Saltzman–Maasch model

Saltzman and Maasch (1990) introduced a global model of the late Cenozoic climate. We will refer to this model as SM90. The model consists of differential equations for the global ice volume x , the atmospheric CO₂ concentration y , and the global mean deep water temperature z , where x , y , and z are dimensionless and denote deviations from 1 Myr averages. Neglecting stochastic perturbations, the model is written in the following deterministic form:

$$\begin{aligned} \tau \dot{x} &= -x - y - vz - uF(t), \\ \tau \dot{y} &= -pz + ry + sz^2 - wyz - z^2y, \\ \tau \dot{z} &= -q(x + z), \end{aligned} \quad (6)$$

where t is the time in kiloyears, and $\tau = 10$ kyr. The parameters p , q , r , s , and w are freely chosen, but the remaining parameters u and v are functions of the other parameters. The insolation forcing $F(t)$ is given by Eq. (5) with $a = 23.58$.

Saltzman and Maasch (1990) simulated glacial cycles for the last 500 kyr by choosing the parameter values as set A in Table 1. Hargreaves and Annan (2002) estimated the parameter values of SM90 as set B using a data assimilation technique.

3.2.2 Maasch–Saltzman model

Prior to the SM90 model, Maasch and Saltzman (1990) studied one of their previous models. This older model, which we call the MS90 model, can be obtained from the SM90 model by setting $v = w = 0$. The other parameter values are chosen as set C in Table 1.

3.2.3 Crucifix–De Saedeleer–Wieczorek model

One of the simplest systems that exhibit self-sustained oscillations is the van der Pol oscillator. With slight modifications and addition of an insolation forcing to the van der Pol oscillator, Crucifix (2012) introduced the following model:

$$\begin{aligned} \tau \dot{x} &= -[y + \beta + \gamma F(t)], \\ \tau \dot{y} &= -\alpha[y^3/3 - y - x], \end{aligned} \quad (7)$$

where x denotes the global ice volume, and y is a conceptual variable for realizing self-sustained oscillations. The insolation forcing $F(t)$ is given by Eq. (5) with $a = 11.77$. For the unforced case $\gamma = 0$, Eq. (7) has one stable equilibrium state for $|\beta| > 1$ and a limit cycle oscillation for $|\beta| < 1$. Following Crucifix (2012), the parameters are set as $\alpha = 30.0$, $\beta = 0.75$, $\gamma = 0.4$, and $\tau = 36.0$ kyr.

The dynamical properties of Eq. (7) were extensively studied by De Saeldeleer, Crucifix, and Wieczorek (2013). Thus, we refer to Eq. (7) as the CSW model.

3.2.4 Paillard–Parrenin model

Paillard and Parrenin (2004) proposed an oscillator model of glacial cycles involving the Antarctic ice-sheet influence on bottom water formation. The model, which we call the PP04 model, consists of equations for the global ice volume x , extent of the Antarctic ice sheet y , and atmospheric CO₂ concentration z :

$$\begin{aligned} \tau_x \dot{x} &= -pz - qF(t) + r - x, \\ \tau_y \dot{y} &= x - y, \\ \tau_z \dot{z} &= \alpha F(t) - \beta x + \gamma H(-w) + \delta - z, \\ w &= ax - by - cI(t) + d, \end{aligned} \quad (8)$$

where $F(t)$ is the summer solstice insolation at 65°N given by Eq. (5) with $a = 23.58$, and $I(t)$ is the insolation at 60°S on February 21. The Heaviside function in the original PP04 model is approximated as $H(x) = 1/2 + \arctan(1000x)/\pi$ for simplicity of the dynamical systems analysis. The parameters are set as $\tau_x = 15$ kyr, $\tau_y = 12$ kyr, $\tau_z = 5$ kyr, $p = 1.3$, $q = 0.5$, $r = 0.8$, $\alpha = 0.15$, $\beta = 0.5$, $\gamma = 0.5$, $\delta = 0.4$, $a = 0.3$, $b = 0.7$, $c = 0$, and $d = 0.27$. The choice $c = 0$ does not change the timing of terminations; it only affects the qualitative agreement with geological records (Paillard and Parrenin 2004).

3.2.5 Imbrie–Imbrie model

Imbrie and Imbrie (1980) proposed a piecewise-linear model for the glacial cycles:

$$\tau \dot{x} = \begin{cases} (1+b)(F(t) - x) & \text{if } F(t) \geq x, \\ (1-b)(F(t) - x) & \text{if } F(t) < x, \end{cases}$$

where x corresponds to a loss in the global ice volume. The variable $F(t)$ is assumed to be the summer solstice insolation at 65°N given by Eq. (5) with $a = 23.58$ (cf. Paillard 2001). The parameters are tuned as $\tau = 17$ kyr and $b = 0.6$. For simplicity of the dynamical systems analysis, we deal with a smoothed system,

$$\tau \dot{x} = [1 + b \tanh k(F(t) - x)](F(t) - x). \quad (9)$$

If the steepness parameter k is large enough, the solution of Eq. (9) converges to that of the original model. We set $k = 10$ in this study.

4 Dynamical systems analysis

4.1 Synchronization under the same astronomical forcing

In what follows, $\mathbf{x}(t)$ represents the climate state at time t for each model described in the previous section². First, we check the dependence of solutions on initial climate states under the same astronomical forcing $F(t)$. We simulate 10^4 solutions $\mathbf{x}^{(i)}(t)$ ($i = 1, 2, \dots, 10^4$), starting from random initial states $\mathbf{x}^{(i)}(t_0)$ at time $t_0 = -20$ Myr, under the same astronomical forcing $F(t)$. For these solutions, we calculate the maximum distance from the average of the ensemble,

$$\Delta_{\max}(t) = \max_{1 \leq i \leq 10^4} |\mathbf{x}^{(i)}(t) - \langle \mathbf{x}^{(i)}(t) \rangle|.$$

We assume that typical solutions converge to the same solution, i.e., *synchronize*, under the same astronomical forcing if $\Delta_{\max}(t) \rightarrow 0$ as $t \rightarrow \infty$. Figure 3 shows the behavior of $\Delta_{\max}(t)$ for each model. In each model except for the SM90-B model, typical solutions synchronize after a certain transient time. This result implies that the attractor is unique at least except for the SM90-B model. It should be mentioned that the transient time is longer than the duration of the Quaternary, approximately 2.58 Myr, in the SM90-A and MS90 models (see Fig. 3).

4.2 Simulated ice volume over the last 700 kyr

Figure 4 shows the simulated normalized ice volume $V(t)$ over the last 700 kyr (solid line) for each model. The $\delta^{18}\text{O}$ record of the LR04 stack is also shown by a dashed line for comparison (Lisiecki and Raymo 2005). In the SM90-A, PP04, and Imbrie models, synchronized solutions after each transient period are well accorded with the $\delta^{18}\text{O}$ record; however, in the MS90 and CSW models, synchronized solutions after each transient period are not accorded with the $\delta^{18}\text{O}$ record, but only some transient solutions agree with the $\delta^{18}\text{O}$ record. In the SM90-B model, solutions starting from different initial states do not synchronize. This result indicates that the SM90-B model exhibits chaotic solutions.

Actually we find coexistence of a quasiperiodic attractor and a chaotic attractor for the SM90-B model.

² For example, $\mathbf{x} = (x, y, z)^T$ for the SM90 model, and $\mathbf{x} = (x)^T$ for the Imbrie model.

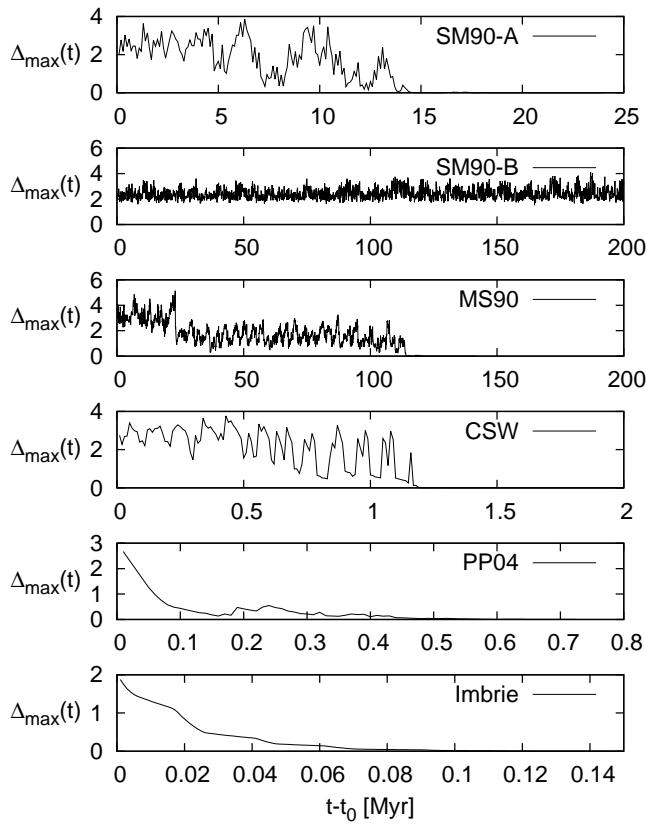


Fig. 3 Convergence of 10^4 solutions for each model: The maximum distance from the average of the ensemble $\Delta_{\max}(t) = \max_{1 \leq i \leq 10^4} |\mathbf{x}^{(i)}(t) - \langle \mathbf{x}^{(i)}(t) \rangle|$ vs. elapsed time $t - t_0$. All solutions start from random initial states at time $t_0 = -20$ Myr. Note that the time scale is different for each panel

The quasiperiodic attractor originates in a stable equilibrium of the unforced system $(x, y, z) = (-0.282, 0.277, 0.282)$. The basin of attraction of this quasiperiodic attractor is much smaller than that of the chaotic attractor, and the amplitude of the quasiperiodic solution is also much smaller than that of the chaotic solutions. Thus, we do not focus on this quasiperiodic solution.

4.3 Types of attractors

We analyze types of attractors using the maximal CLE λ , the phase sensitivity exponent μ , and the parameter sensitivity exponent ν . As shown in Table 2, the maximal CLE λ is positive for the SM90-B model and is negative for all the other models. Thus, the SM90-B model is chaotic, and all the other models are non-

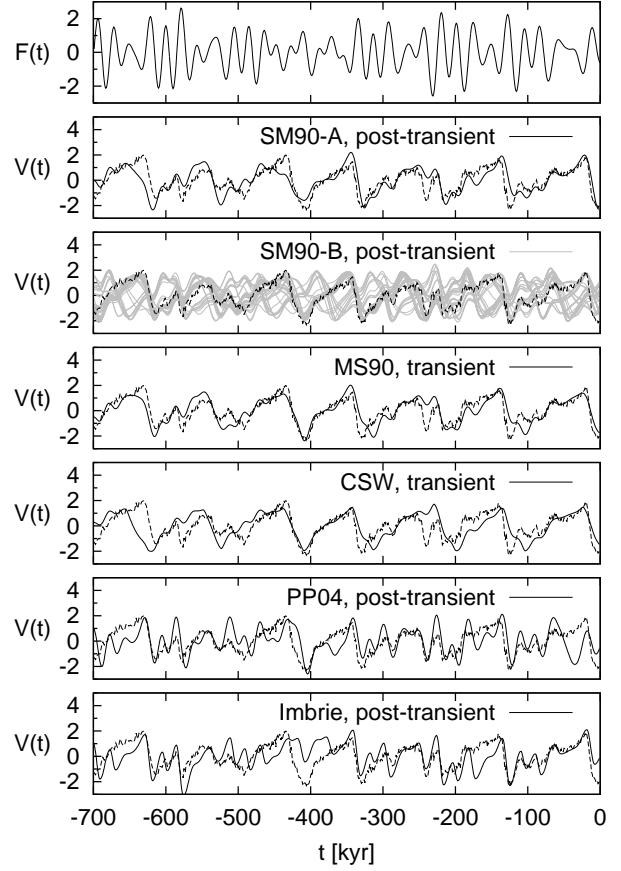


Fig. 4 Top panel: Normalized insolation $F(t)$ at 65°N on the day of summer solstice. Other panels: Simulated normalized ice volume $V(t)$ (solid lines) for the models described in Section 3. The $\delta^{18}\text{O}$ record of the LR04 stack is also shown for comparison (dashed lines). Post-transient solutions are considered for the SM90-A, SM90-B, PP04, and Imbrie models, and transient solutions are considered for the MS90 and CSW models

Table 2 Characteristic exponents and types of attractors for the different models. The asterisk (*) indicates that the sensitivity exponents are undefined for the SM90-B model.

Models	$\lambda [\text{kyr}^{-1}]$	μ	ν	Attractor
SM90-A	-0.0014	1.06	1.15	SNA
SM90-B	+0.0020	*	*	Chaotic
MS90	-0.0005	0.95	1.09	SNA
CSW	-0.016	1.10	1.13	SNA
PP04	-0.0093	1.09	1.08	SNA
Imbrie	-0.046	0.0	0.0	Quasiperiodic

chaotic. This result is consistent with that of convergence experiment.

The next question is whether the attractors are strange or not. Figure 5(a) shows the phase sensitivity functions $\Gamma(t)$ of x with respect to a phase variable θ_4 for the models with nonchaotic attractors, and Fig. 5(b) shows

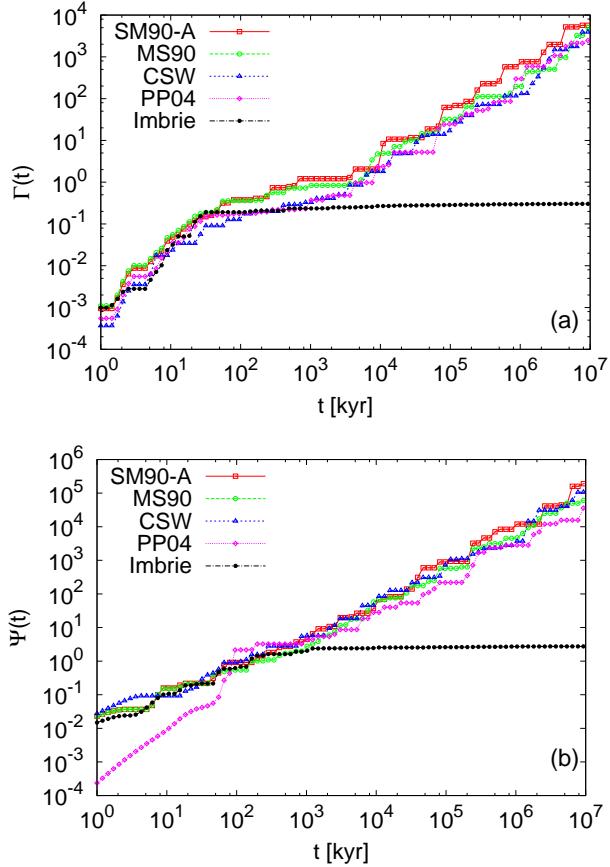


Fig. 5 **a** Phase sensitivity function $\Gamma(t)$ of x with respect to a phase variable θ_4 . **b** Parameter sensitivity function $\Psi(t)$ of x with respect to u for the SM90-A and MS90 models, parameter β for the CSW model, and parameter b for the Imbrie model

the parameter sensitivity functions $\Psi(t)$ of x with respect to parameter u for the SM90-A and MS90 models, parameter β for the CSW model, and parameter b for the Imbrie model. Here, the functions $\Gamma(t)$ and $\Psi(t)$ are calculated for a set of 20 initial conditions. For the Imbrie model, the functions $\Gamma(t)$ and $\Psi(t)$ saturate as time increases, but for the other models, the functions $\Gamma(t)$ and $\Psi(t)$ grow in power-laws with exponents $\mu > 0$ and $\nu > 0$, respectively. Table 2 shows the values of μ and ν , which are estimated from the power-law behaviors for $t > 10^4$. Therefore, the attractors of the SM90-A, MS90, CSW, and PP04 models are SNAs, and the attractor of the Imbrie model is a quasiperiodic attractor.

4.4 Illustration of attractors by simplification of astronomical forcing

Since the astronomical forcing $F(t)$ in Eq. (5) has high degrees of freedom, namely, $N = 35$ in our study, the

attractors in climate models have at least 35 dimensions, which are too high-dimensional to be visualized. To capture the geometric features of these attractors, we illustrate approximated attractors by a truncation of the astronomical forcing. Leaving only the 23 kyr ($i = 1$) and 19 kyr ($i = 3$) precession terms and the 41 kyr ($i = 4$) obliquity term, we simplify the astronomical forcing $F(t)$ as

$$F(t) = \frac{1}{a_{\{1,3,4\}}} \sum_{i \in \{1,3,4\}} [s_i \sin(\omega_i t) + c_i \cos(\omega_i t)].$$

These three frequency components constitute 78% of the original astronomical forcing with respect to standard deviations. The scaling constant is $a_{\{1,3,4\}} = 9.08$ for the CSW model and $a_{\{1,3,4\}} = 18.3$ for the other models.

In spite of the simplification of the astronomical forcing, we obtain the same type of attractor for each model although the values of characteristic exponents are slightly changed: an SNA with $\lambda \approx -0.0004$, $\mu = 1.39$, and $\nu = 1.43$ for the SM90-A model, a chaotic attractor with $\lambda \approx 0.0025$ for the SM90-B model, an SNA with $\lambda \approx -0.00012$, $\mu = 1.09$, and $\nu = 1.35$ for the MS90 model, an SNA with $\lambda \approx -0.0088$, $\mu = 1.13$, and $\nu = 1.10$ for the CSW model, an SNA with $\lambda \approx -0.0089$, $\mu = 1.07$, and $\nu = 0.94$ for the PP04 model, and a quasiperiodic attractor with $\lambda \approx -0.046$ and $\nu_\theta = \nu_u = 0$ for the Imbrie model.

The stroboscopic map is useful for visualizing attractors having three or more degrees of freedom. Observing the system at the moments $t_n = 2\pi n/\omega_1$ ($n \in \mathbb{Z}$), where a phase $\theta_1(t)$ attains a constant value, the system reduces to a stroboscopic map $(\theta_3(t_n), \theta_4(t_n), \mathbf{x}(t_n)) \mapsto (\theta_3(t_{n+1}), \theta_4(t_{n+1}), \mathbf{x}(t_{n+1}))$. Figure 6 shows the attractors of the simplified models. Each panel presents the projection of the attractor on the space $(\theta_3(t_n)/2\pi, \theta_4(t_n)/2\pi, x(t_n))$. The SM90-A, MS90, CSW, and PP04 models exhibit SNAs, and the dependence of the climate state on the phase of astronomical forcing is not given by smooth relations, but constitutes a geometrically strange set. Moreover, the dependence seems to be discontinuous almost everywhere. For the Imbrie model, the attractor is quasiperiodic, and the dependence is given by a smooth graph. For the SM90-B model, the attractor is chaotic, and the dependence constitutes a geometrically strange set, as in the case of SNAs.

5 A dynamical property of SNAs: Sensitivity to perturbations

If a climate model written as Eq. (1) has a chaotic attractor, solutions sensitively depend on initial climate

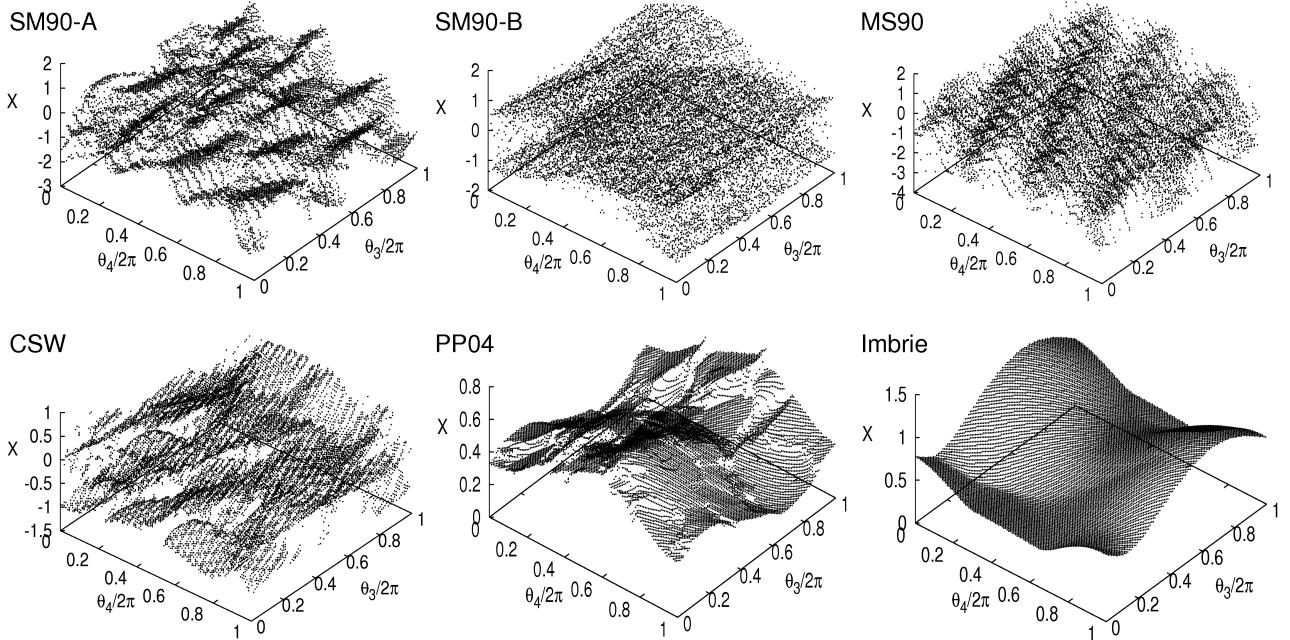


Fig. 6 Attractors of each model for a simplified astronomical forcing. Each panel shows the projection of the attractor on the space $(\theta_3/2\pi, \theta_4/2\pi, x)$

states $\mathbf{x}(t_0)$. On the other hand, if a climate model has an SNA or a quasiperiodic attractor, we obtain one or several synchronized solutions after a certain transient time. Then what is the difference between SNAs and quasiperiodic attractors?

Khovanov et al. (2000) stated that SNAs and quasiperiodic attractors differ in sensitivity to perturbations. We illustrate the difference by simulating the CSW and Imbrie models in the presence of noise. In the former, we add independent white Gaussian noises $\xi_i(t)$ ($i = x$ or y) to the right-hand sides of Eq. (7), respectively, where $\langle \xi_i(t) \rangle = 0$, $\langle \xi_i(t) \xi_j(t') \rangle = \sigma^2 \delta_{ij} \delta(t - t')$, and $\sigma = 0.1 \text{ kyr}^{1/2}$. In the latter, we add a white Gaussian noise $\xi_x(t)$ with the same average and standard deviation to the right-hand side of Eq. (9).

Figure 7(a) shows two trajectories of the CSW model simulated with and without noises, and Fig. 7(d) shows those of the Imbrie model. The CSW model is more sensitive to noise than the Imbrie model because of the local instability of SNA. To characterize the local instability of attractors, let us define a finite-time CLE

$$\lambda_T(t) = \frac{1}{T} \ln \frac{|\delta \mathbf{x}(t + T)|}{|\delta \mathbf{x}(t)|}.$$

The finite-time CLE $\lambda_T(t)$ gives the rate of exponential divergence of nearby orbits during the time interval $[t, t + T]$. Note that $\lambda_T(t)$ converges to the maximal CLE λ as $T \rightarrow \infty$. For SNAs, the finite-time CLE $\lambda_T(t)$

can be positive even for large t with nonzero probability (Pikovsky and Feudel 1995). Thus, the orbits of SNAs are sensitive to perturbations in time intervals with $\lambda_T(t) > 0$. Figures 7(a)–(c) show that the perturbed trajectory occasionally deviates after the finite-time CLE $\lambda_T(t)$ becomes positive, where $T = 70 \text{ kyr}$. This local instability may be attributed to a repellor embedded in SNA as mentioned in Section 2. Such a temporary loss of synchronization has been shown also in another model of glacial cycles (Crucifix 2011).

On the other hand, for quasiperiodic attractors, the finite-time CLE $\lambda_T(t)$ takes only negative values for large T . As a result, the orbits of quasiperiodic attractors are robust to perturbations in comparison with those of SNAs. Figures 7(d)–(f) show that the Imbrie model is robust to the same magnitude of noise that can cause deviations of trajectories in the CSW model.

6 Conclusion

We analyzed five conceptual models of glacial cycles in the framework of the dynamical systems theory. The SM90-A, MS90, CSW, and PP04 models exhibit a strange nonchaotic attractor (SNA), and the dependence of the climate state on the phase of astronomical forcing is not given by smooth relations, but constitutes a geometrically strange set. The SM90-B model exhibits a chaotic attractor, and the dependence of the climate

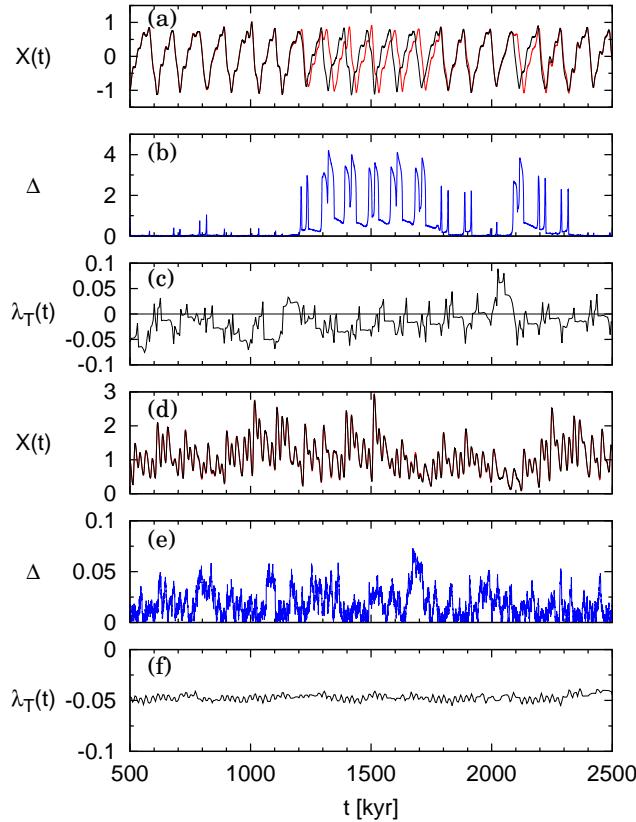


Fig. 7 Sensitivity of climate models to noise: **a** Ice volume $x(t)$ simulated by the CSW model with noise (black) and without noise (red). **b** Distance $\Delta(t)$ in the space (x, y) between perturbed and unperturbed trajectories corresponding to panel **a**. **c** The finite-time CLE $\lambda_T(t)$ of the unperturbed trajectory in panel **a**. $T = 70$ kyr. **d** Loss of ice volume $x(t)$ simulated by the Imbrie model with noise (black) and without noise (red). **e** Distance $\Delta(t)$ between perturbed and unperturbed trajectories of panel **d**. **f** The finite-time CLE $\lambda_T(t)$ of the unperturbed trajectory in panel **d**. $T = 70$ kyr

state on the phase of astronomical forcing is not given by smooth relations, but constitutes a geometrically strange set. The Imbrie model exhibits a quasiperiodic attractor, and the dependence of the climate state on the phase of astronomical forcing is given by a smooth graph. These results suggest that the dynamics of SNAs is a candidate for the dynamics of glacial cycles, in addition to the well-known dynamics of quasiperiodicity and chaos³. Since SNAs have repellors in themselves, the orbits on SNAs are locally unstable. Therefore, if the real climate system has an SNA, the orbit will be temporarily sensitive to perturbations when the finite-time conditional Lyapunov exponent is positive.

Brindley and Kapitaniak (1992) specified an *inhibition of chaotic behavior* due to the quasiperiodicity of

³ Chaos in conceptual models of glacial cycles has often been mentioned in the literature (e.g. Nicolis 1987; Saltzman and Verbitsky 1992, 1993; Ghil 1994; Huybers 2009).

forcing. In their article, the inhibition of chaotic behavior refers to the appearance of statistical periodicity for the ensemble of orbits, and it is different from the non-chaoticity based on nonpositive Lyapunov exponents in this study.

The analysis in this paper was restricted to spatially zero-dimensional models that can be approximated by smooth dynamical systems. A *hybrid dynamical system* is a system that has both continuous and discrete dynamics (Aihara and Suzuki 2010). A number of ice age models are expressed as hybrid dynamical systems, more precisely, *hybrid automata*⁴ (van der Schaft and Schumacher 2000). For example, the models by Paillard (1998), Parrenin and Paillard (2003), and Ashkenazy and Tziperman (2004) are hybrid automata. The Imbrie and PP04 models are also hybrid dynamical systems but can be approximated by smooth dynamical systems because they are especially *piecewise affine systems*. The description of the dynamics appearing in hybrid automata models or in more complicated models with spacial dimensions is still unclear. This problem will be studied in the future.

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Appendix 1: Parameter values for the insolation formula

The parameter values for the insolation formula (5) given by De Saedeleer et al. (2013) are listed in Table 3 to make the paper self-contained. The frequencies and coefficients are arranged in descending order of their power $s_i^2 + c_i^2$.

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⁴ A simple example of hybrid automata is the temperature-thermostat system, where a continuous variable is the temperature and a discrete variable is the status of the heater {on, off}.

Table 3 Parameter values for the insolation formula of De Saedeleer et al. (2013)

Index i	ω_i [rad/kyr]	s_i [W/m ²]	c_i [W/m ²]
1	0.264933601588513	-15.549049332290400	-9.704062871105320
2	0.280151350350945	15.431955636170100	4.752472711315250
3	0.331110950251899	9.099224935273400	-10.611524488739001
4	0.153249478547167	-11.228737681512399	3.516820752112410
5	0.328024059125949	-7.870653840136690	6.615442460635030
6	0.326211762560183	0.813786144754451	-4.526414080992460
7	0.158148666238883	-3.824993714675400	-0.761851750263805
8	0.269742342439881	0.069044850431486	-3.316392609695580
9	0.117190147169570	2.288148059560660	1.802337026846230
10	0.332923246817665	1.440507707859670	1.063392860501200
11	0.150162587421217	1.549041763533020	-0.088394191276982
12	0.217333905941751	0.380973541305497	-1.463017119992100
13	0.155061775112933	-1.297700819564400	-0.635152963728496
14	0.371638925683567	0.925324276580528	-1.020667586721540
15	0.324937167999999	-0.576082669762308	1.186695727393380
16	0.275366617473065	0.997628846513796	-0.362906496840039
17	0.211709630908568	-0.810768209286259	-0.577980646565494
18	0.156336369673117	-0.918358442095885	0.196083726889428
19	0.334197841377850	0.346906064369828	-0.648189701487285
20	0.259396912994958	0.339477750517033	-0.560509461538342
21	0.323124871434233	-0.378637986107629	0.527217891742183
22	0.148350290855451	0.256895610735773	-0.524697312305024
23	0.111684123041346	-0.428006728186239	0.357006342316690
24	0.336010137943616	-0.421809264016129	0.324327509437558
25	0.274551083291250	-0.441772417569753	0.289576210423804
26	0.126901871803777	0.257509918217341	-0.377639794223366
27	0.177861471704732	-0.161827722328271	-0.362683869407858
28	0.206924898030688	-0.335783913402678	-0.019479215012864
29	0.212525165090383	0.267659228540196	0.128915417116900
30	0.004899187691716	-0.16047984721994	0.059407796893426
31	0.418183080135680	-0.018488406464501	0.109632390175297
32	0.433400828898112	-0.004919921945456	-0.106148873639336
33	0.229992875969202	0.069618973318896	0.074623171406128
34	0.311398144786051	0.013835372762118	0.030473666884042
35	0.306498957094334	0.024734974816962	0.014046439534097

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